

# Observational constraints on the physical nature of submillimetre source multiplicity: chance projections are common

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## ABSTRACT

Interferometric observations have demonstrated that a significant fraction of single-dish submillimetre (submm) sources are blends of multiple submm galaxies (SMGs), but the nature of this multiplicity, i.e. whether the galaxies are physically associated or chance projections, has not been determined. We performed spectroscopy of 11 SMGs in six multicomponent submm sources, obtaining spectroscopic redshifts for nine of them. For an additional two component SMGs, we detected continuum emission but no obvious features. We supplement our observed sources with four single-dish submm sources from the literature. This sample allows us to statistically constrain the physical nature of single-dish submm source multiplicity for the first time. In three (3/7, or  $43^{+39}_{-33}$  per cent at 95 per cent confidence) of the single-dish sources for which the nature of the blending is unambiguous, the components for which spectroscopic redshifts are available are physically associated, whereas 4/7 ( $57^{+33}_{-39}$  per cent) have at least one unassociated component. When components whose spectra exhibit continuum but no features and for which the photometric redshift is significantly different from the spectroscopic redshift of the other component are also considered, 6/9 ( $67^{+26}_{-37}$  per cent) of the single-dish sources are comprised of at least one unassociated component SMG. The nature of the multiplicity of one single-dish source is ambiguous. We conclude that physically associated systems and chance projections both contribute to the multicomponent single-dish submm source population. This

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result contradicts the conventional wisdom that bright submm sources are solely a result of merger-induced starbursts, as blending of unassociated galaxies is also important.

**Key words:** galaxies: high-redshift – galaxies: starburst – infrared: galaxies – submillimetre: galaxies.

## 1 INTRODUCTION

Submillimetre (submm) wavelengths are ideal for selecting galaxies across a wide range of redshift because of the so-called negative  $k$ -correction (e.g. Blain et al. 2002), but their drawback is that the beam sizes of typical single-dish submm telescopes are  $\gtrsim 15$  arcsec ( $\gtrsim 120$  kpc at  $z \gtrsim 1$ ).<sup>1</sup> Consequently, blending of emission from more than one galaxy into one single-dish submm source is possible. Such blending is more likely for submm sources than for e.g. optical sources because of both the large beam size and the negative  $k$ -correction, which jointly imply, roughly speaking, that a galaxy with a fixed spectral energy distribution will contribute approximately the same flux to the observed single-dish submm source if the galaxy is located anywhere within a cylinder of diameter  $\sim 240$  kpc spanning  $z \sim 1$ –10. This potential problem has long been recognized (e.g. Hughes et al. 1998), but constraining the prevalence and nature of blended submm sources has been challenging. The fact that single-dish submm sources often have multiple radio (e.g. Ivison et al. 2002, 2007) and  $K$ -band (Smith et al. 2017) counterparts suggests that blending is common. However, to directly determine whether single-dish submm sources are actually blends of multiple sources, interferometric observations of the dust continuum emission are required.

Before the Atacama Large Millimeter/submillimeter Array (ALMA<sup>2</sup>) came online, only a handful of single-dish submm sources were observed with submm interferometers such as the Submillimeter Array (SMA<sup>3</sup>) and the Plateau de Bure Interferometer<sup>4</sup>; some of these pre-ALMA interferometric studies resolved single-dish submm sources into multiple components (e.g. Younger et al. 2009; Engel et al. 2010; Wang et al. 2011; Barger et al. 2012; Smolčić et al. 2012). In the era of ALMA, it has become possible to interferometrically map large numbers of single-dish submm sources. To date, ALMA follow-up observations have demonstrated that many single-dish submm sources are blends of two or more resolved SMGs (e.g. Karim et al. 2013; Hodge et al. 2013; Wiklind et al. 2014; Bussmann et al. 2015; Simpson et al. 2015a,b; see Casey, Narayanan & Cooray 2014 for a review), although details such as the fraction of single-dish submm sources that are blends and the effect of blending on submm number counts are still debated (e.g. Chen et al. 2013; Koprowski et al. 2014; Michałowski et al. 2017).

Although it is clear that a significant fraction of single-dish submm sources are blends, the physical nature of this multiplicity has not been constrained. Specifically, are the individual components of multicomponent single-dish submm sources physically associated galaxies, either undergoing a merger or within the same dark matter structure (e.g. group or filament) but not actively merging? Or are blended submm sources chance projections of galaxies at different redshifts that have no dynamical influence on one another?

Various theoretical works have considered the effects of blending on submm number counts. Hayward et al. (2011) were the first to suggest that pre-coalescence galaxy mergers (i.e. blends of two physically associated SMGs) should contribute significantly to the single-dish submm number counts. In follow-up work, Hayward et al. (2012) discussed how to distinguish pre-coalescence mergers (galaxy-pair SMGs) from coalescence-stage starbursts, and Hayward et al. (2013a) presented detailed predictions for the relative contributions of merger-induced starbursts, blended pre-coalescence merging galaxies, and isolated discs to the submm number counts. In the model of Narayanan et al. (2015), the contribution of satellite galaxies results in blended sources comprised of physically associated component SMGs. Hayward et al. (2013b) were the first to investigate the effects of blending of *unassociated* galaxies; they predicted that of the subset of single-dish submm sources that are blends of multiple SMGs, the majority should be comprised of at least one SMG that is physically unassociated with the other component(s). Using a semi-analytic model in which the physical nature of SMGs is drastically different than in the Hayward et al. (2013b) model, Cowley et al. (2015) concluded that effectively all multicomponent single-dish submm sources should be comprised of at least one unassociated SMG. Applying a ‘counts-matching’ approach to a semi-analytic model, Muñoz Arancibia et al. (2015) also predicted that the majority of the components of blended single-dish submm sources are spatially unassociated. Notably, all theoretical models that treat blending of unassociated galaxies predict that chance projections contribute significantly to the multiple-component single-dish submm source population, but these predictions have not been tested to date.

It is possible to indirectly constrain the relative importance of the two types of blending using the distribution of the angular separations of the resolved submm components of single-dish submm sources, but such analyses have yielded conflicting results (see the comparison presented in fig. 7 of Bussmann et al. 2015), and the indirect nature of such constraints makes redshift-based constraints preferable. Interferometric maps of the dust continuum emission are a prerequisite for obtaining accurate redshifts because relying on potential counterparts at other wavelengths runs the risk of yielding a redshift for a galaxy that does not actually contribute to the single-dish submm source. Even when the positions of the resolved submm components are known, obtaining redshifts is challenging because photometric redshifts of SMGs may not be sufficiently accurate to constrain the physical nature of the blending (Simpson et al. 2014). Spectroscopic redshifts can provide unambiguous constraints, but obtaining spectroscopic redshifts for SMGs is notoriously difficult for multiple reasons. For example, at fixed submm flux density, the optical/near-infrared (NIR) emission line luminosities of individual SMGs can differ considerably, probably owing to patchy dust extinction. Consequently, optical/NIR spectroscopic follow-up often yields non-detections (e.g. Danielson et al. 2017). Obtaining redshifts via molecular and atomic emission lines in the far-infrared (FIR) and submm is an alternative approach, and even unresolved observations can reveal chance projections (Zavala et al. 2015). However, the overhead associated with multiple tunings makes this observationally expensive, and often, guided by photometric

<sup>1</sup> Throughout this work, we assume  $\Omega_m = 0.31$ ,  $\Omega_\Lambda = 0.69$ , and  $H_0 = 68 \text{ km s}^{-1} \text{ Mpc}^{-1}$  (Planck Collaboration XIII 2016).

<sup>2</sup> <http://www.almaobservatory.org/>

<sup>3</sup> <https://www.cfa.harvard.edu/sma/>

<sup>4</sup> [www.iram-institute.org/EN/plateau-de-bure.php](http://www.iram-institute.org/EN/plateau-de-bure.php)

redshifts, only a limited wavelength – and thus redshift – range is probed.<sup>5</sup> For these reasons, very few multicomponent single-dish submm sources for which spectroscopic redshifts of at least two of the component SMGs are available have been presented in the literature (Wang et al. 2011; Barger et al. 2012; Danielson et al. 2017).

We present spectroscopic constraints on the redshifts of components of six single-dish submm sources based on spectra obtained with Keck,<sup>6</sup> Gemini,<sup>7</sup> and the Very Large Telescope.<sup>8</sup> We supplement our observational data set with four single-dish submm sources from the literature. The combined sample is comprised of seven multicomponent single-dish submm sources for which optical emission line-based spectroscopic redshifts are available for two or more components. For an additional three single-dish sources, a spectroscopic redshift is available for one component, and continuum only is detected in the spectrum of the other; for these sources, photometric redshift estimates for the continuum-only components are available, and we use these to tentatively constrain the nature of the multiplicity of these sources. For the first time, we are able to constrain the relative contributions of physically associated galaxies and chance projections to the multicomponent single-dish submm source population.

## 2 OBSERVATIONAL DATA

### 2.1 S2CLS sources observed in this study

Details regarding our single-dish submm source sample, including both those observed in this work and those drawn from the literature, are presented in Table 1. The single-dish submm sources observed as part of this study were primarily drawn from the Submillimetre Common-User Bolometer Array 2 (SCUBA-2<sup>9</sup>) Cosmology Legacy Survey at 850  $\mu\text{m}$  (S2CLS;<sup>10</sup> Geach et al. 2017). Simpson et al. (2015a,b) followed up 30 bright ( $S_{870} \gtrsim 8$  mJy) S2CLS single-dish submm sources with ALMA. An additional 75 single-dish submm sources with  $S_{870} > 8$  mJy were followed up with the SMA; the catalogue and SMA observations are detailed in Hill et al. (2017). The interferometric observations reveal whether the S2CLS single-dish submm sources are composed of multiple SMGs. We selected the brightest multiple-component single-dish submm sources from the S2CLS for which SMA or ALMA follow-up observations were available prior to our observing nights as candidates for NIR spectroscopic follow-up. To avoid selecting an incorrect counterpart while simultaneously minimizing the chance of following up spurious ALMA or SMA sources, we targeted multiples for which at least one – and preferably more than one – of the ALMA or SMA sources had an unambiguous companion at optical or NIR wavelengths. Note that this selection may make our sample biased towards lower redshift or/and more-massive SMGs and against chance projections in which the secondary submm compo-

nent is located at a significantly greater redshift than the primary and thus may not be visible in the rest-frame optical.

#### 2.1.1 MOSFIRE observations

We observed two of the component SMGs of each of two multicomponent single-dish SCUBA-2 850  $\mu\text{m}$  sources (LOCK-03 and LOCK-09) using the Multi-Object Spectrometer for InfraRed Exploration (MOSFIRE;<sup>11</sup> McLean et al. 2010, 2012), an NIR imaging spectrometer on the Keck 1 telescope, on 2016 February 25. We additionally observed the two component SMGs of LOCK-08; because these are separated by 25 arcsec, they are not blended in the SCUBA-2 map, but they would be blended in e.g. LABOCA 870  $\mu\text{m}$  or SPIRE 350 and 500  $\mu\text{m}$  maps. We detected line or/and continuum from all six of the resolved SMGs. The slitmasks were designed to also include some single-component submm sources, only one component of some multicomponent submm sources, and any nearby radio sources that could be accommodated within the 3 arcmin  $\times$  6.1 arcmin field of view.

All slits were 0.7 arcsec in width, resulting in spectra with resolving power  $R \approx 3650$  in the *H*- or *K*-band atmospheric windows. Given that the typical sizes of SMGs ( $R_c \sim 1$  kpc, as determined from submm emission; e.g. Simpson et al. 2015b) are much smaller than the slit width (approximately 6 kpc at the redshifts of interest) and that the seeing was good (0.65 arcsec), slit losses are likely  $\lesssim 30$  per cent. Moreover, because we are primarily concerned with redshifts rather than line luminosities, aperture corrections to the line fluxes are unnecessary.

The spectra were obtained using the standard two-position ‘mask nod’ in which the telescope position was dithered  $\pm 1.5$  arcsec along the slit direction, with individual integration times of 120 s (*H* band) or 180 s (*K* band). Total integration times of 2880 s were obtained, with 24 and 16 individual exposures in the *H* and *K* bands, respectively.

The MOSFIRE data were reduced using the publicly available data reduction pipeline developed by the instrument team;<sup>12</sup> see Steidel et al. (2014) for a detailed description of the procedure. One-dimensional spectra were extracted, flux-calibrated, and analysed using the MOSPEC package (Strom et al. 2017).

#### 2.1.2 GNIRS observations

Near-infrared spectra of resolved components of two multicomponent single-dish submm sources (UDS292 and UDS306) were obtained using the cross-dispersed mode of the Gemini Near-Infrared Spectrograph (GNIRS<sup>13</sup>) on the Gemini North 8.1 m telescope. This configuration provides continuous spectral coverage from 0.84 to 2.48  $\mu\text{m}$  at a spectral resolution of  $R \approx 1500$  with a spatial scale of 0.15 arcsec  $\text{pix}^{-1}$ . The slit dimensions were 1.0 arcsec  $\times$  7.0 arcsec. The observations used an ABBA pattern, nodding along the slit to keep the galaxy on slit at all times. Eight individual on-source integrations of 240 s each were performed for each source. Observations of standard stars were obtained before and after each set of SMG observations. These were used to correct the spectra for telluric absorption.

The spectral reduction, extraction, and wavelength and flux calibration procedures were performed using the Gemini IRAF package

<sup>5</sup> Given the bias against chance projections inherent in these observations, in this work, we do not consider FIR/submm atomic or molecular gas emission line-based redshifts from the literature except for COSBO3, for which we complement  $H\alpha$ -based redshifts with some CO-based redshifts. Excluding the CO-based redshifts does not qualitatively affect our conclusions.

<sup>6</sup> <http://www.keckobservatory.org/>

<sup>7</sup> <http://www.gemini.edu/>

<sup>8</sup> <http://www.eso.org/public/usa/teles-instr/paranal-observatory/vlt/>

<sup>9</sup> <http://www.eaobservatory.org/jcm/instrumentation/continuum/scuba-2/>

<sup>10</sup> <http://www.astro.dur.ac.uk/irs/S2CLS/>

<sup>11</sup> <https://www2.keck.hawaii.edu/inst/mosfire/>

<sup>12</sup> <https://keck-datareductionpipelines.github.io/MosfireDRP/>

<sup>13</sup> <http://www.gemini.edu/sciops/instruments/gnirs/>

**Table 1.** Properties of the components of the single-dish submm sources in our sample.

Single-dish ID	Component ID	RA (J2000.0) (h:m:s)	Dec. (J2000.0) (d:m:s)	$S_{860/870}^a$ (mJy)	$z_{\text{spec}}^b$	$z_{\text{phot}}$	Lines detected	Submm ref <sup>c</sup>	$z$ ref <sup>d</sup>	Figure
LOCK-03	a	10:47:27.97	+58:52:14.1	$8.1 \pm 1.8$	2.209	$2.76 \pm 0.18$	H $\alpha$ , [N II]	1	This work [M] (7)	2
	b	10:47:26.52	+58:52:12.8	$8.0 \pm 1.9$	2.363	$2.10 \pm 0.34$	H $\alpha$ , [N II]	1	This work [M] (7)	2
LOCK-08	a	10:47:00.18	+59:01:07.5	$10.4 \pm 1.6$	2.279	$2.06^{+0.09}_{-0.18}$	H $\alpha$	1	This work [M] (8)	2
	b	10:47:02.48	+59:00:50.3	$4.8 \pm 1.6$	2.280	–	H $\alpha$ , [N II]	1	This work [M]	2
LOCK-09	a	10:45:23.11	+59:16:18.6	$9.4 \pm 1.5$	–	–	–	1	–	–
	b	10:45:24.94	+59:16:26.7	$5.1 \pm 1.5$	1.633	$1.46 \pm 0.19$	H $\alpha$ , [N II]	1	This work [M] (7)	1
	c	10:45:23.55	+59:16:32.2	$4.5 \pm 1.5$	<sup>e</sup>	$0.90 \pm 0.05$	Continuum	1	This work [M] (7)	1
UDS292	0	2:17:21.53	–5:19:07.8	$4.2 \pm 0.8$	2.383	$2.65^{+0.25}_{-0.07}$	[O III]4959, 5007	2	This work [G] (9)	3
	1	2:17:21.96	–5:19:09.8	$3.9 \pm 0.8$	2.387	$2.51^{+0.23}_{-0.10}$	H $\alpha$	2	This work [G] (9)	3
UDS306	0	2:17:17.07	–5:33:26.6	$8.3 \pm 0.5$	2.603	$2.31^{+0.06}_{-0.21}$	H $\alpha$ , [O II]	2	This work [G] (9)	3
	1	2:17:17.16	–5:33:32.5	$2.6 \pm 0.4$	<sup>e</sup>	$1.28^{+0.53}_{-0.06}$	Continuum	2	This work [G] (9)	3
	2	2:17:16.81	–5:33:31.8	$3.0 \pm 0.9$	2.606 <sup>f</sup>	–	[O III]4959, 5007	2	This work [X]	4
COSBO3	a	10:00:56.95	+2:20:17.3	$5.3 \pm 0.3$	2.494	–	CO(1–0), CO(3–2)	3	10, 11	–
	b	10:00:57.57	+2:20:11.2	$3.8 \pm 0.3$	2.513	$2.71^{+0.15}_{-0.13}$	H $\alpha$ , CO(1–0)	3	10, 11 (12)	–
	c	10:00:57.27	+2:20:12.7	$1.7 \pm 0.2$	2.498	$2.18^{+0.10}_{-0.08}$	H $\alpha$	3	This work [M] (12)	2
	d	10:00:57.40	+2:20:10.8	$2.2 \pm 0.4$	2.508	$2.31^{+0.03}_{-0.04}$	CO(1–0)	3	10 (12)	–
	e	10:00:56.86	+2:20:08.9	$1.69 \pm 0.3$	2.503	$2.28^{+0.05}_{-0.06}$	CO(1–0)	3	10 (12)	–
GOODS 850-13	a	12:37:14.03	+62:11:56.4	$3.2 \pm 0.9$	–	$3.46^{+1.04}_{-0.86}$	–	4	(3)	–
	b	12:37:14.26	+62:12:08.1	$4.1 \pm 0.7$	3.157	1.25	Not reported	4	17 (18)	–
	c	12:37:12.00	+62:12:12.3	$5.3 \pm 0.9$	2.914	2.80	Ly $\alpha$ , C IV	4	19 (16)	–
GOODS 850-15	a	12:36:21.10	+62:17:09.6	$3.4 \pm 0.6$	1.988	2.91	H $\alpha$ , interstellar abs. lines	5	19 (17)	–
	b	12:36:21.30	+62:17:08.1	$3.5 \pm 0.7$	1.992	2.016	H $\alpha$ , [N II]	5	18 (19)	–
ALESS 084	1	3:31:54.50	–27:51:05.6	$3.2 \pm 0.6$	3.965	$1.92^{+0.09}_{-0.07}$	Ly $\alpha$ , N v, cont. ( $Q = 3$ )	6	20 (21)	–
	2	3:31:53.85	–27:51:04.4	$3.2 \pm 0.8$	<sup>e</sup>	$1.75^{+0.08}_{-0.19}$	cont., poss. faint lines ( $Q = 4$ )	6	20 (21)	–
ALESS 088	1	3:31:54.76	–27:53:41.5	$4.6 \pm 0.6$	1.268	$1.84^{+0.12}_{-0.11}$	[O II], [O II]3726,3729 ( $Q = 1$ )	6	20 (21)	–
	2	3:31:55.39	–27:53:40.3	$2.1 \pm 0.5$	2.519	–	C III]2326, C IV ( $Q = 3$ )	6	20	–
	5	3:31:55.81	–27:53:47.2	$2.9 \pm 0.7$	2.294	$2.30^{+0.11}_{-0.50}$	Ly $\alpha$ , He II, cont. ( $Q = 2$ )	6	20 (21)	–
	11	3:31:54.95	–27:53:37.6	$2.5 \pm 0.7$	2.358	$2.57^{+0.04}_{-0.12}$	Ly $\alpha$ , C III] ( $Q = 3$ )	6	20 (21)	–

<sup>a</sup>Peak SMA 860  $\mu\text{m}$  (Hill et al. 2017) or ALMA 870  $\mu\text{m}$  (Simpson et al. 2015b) flux density.

<sup>b</sup>The errors on the spectroscopic redshifts reported in this work are  $dz \sim 0.0005$ .

<sup>c</sup>Source of submm interferometric observations: (1) SMA (Hill et al. 2017); (2) ALMA (Simpson et al. 2015a,b); (3) ALMA (Bussmann et al. 2015); (4) SMA (Wang et al. 2011); (5) SMA (Barger et al. 2012); (6) ALMA (Hodge et al. 2013).

<sup>d</sup>Reference for  $z_{\text{spec}}$  ( $z_{\text{phot}}$  reference in parentheses) or spectral follow-up if undetected or continuum-only detection: (7) Strazzullo et al. (2010); (8) Roseboom et al. (2012); (9) Simpson et al. (2017); (10) Wang et al. (2016); (11) Champagne et al. in preparation; (12) Laigle et al. (2016); (13) Barger, Cowie & Wang (2008); (14) Trouille et al. (2008); (15) Chapman et al. (2005); (16) Bluck et al. (2012); (17) Wang, Cowie & Barger (2006); (18) Swinbank et al. (2004); (19) Barger et al. (2014); (20) Danielson et al. (2017) – the  $Q$  values represent the spectral quality, with  $Q = 3$  indicating redshifts that should not be considered fully independent of the photometric redshifts and  $Q = 4$  corresponding spectra that exhibit continuum but no sufficiently trustworthy features to assign a spectroscopic redshift; see Danielson et al. (2017) for further details; (21) Simpson et al. (2014). For spectroscopic redshifts obtained in this work, the instrument is denoted by a letter in brackets: M = MOSFIRE, G = GNIRS, X = XSHOOTER.

<sup>e</sup>Spectrum exhibits continuum only; thus, the spectroscopic redshift is unconstrained.

<sup>f</sup>The NIR source is offset by 2.5 arcsec ( $\sim 20$  kpc at  $z = 2.6$ ) from the ALMA position. It is thus likely that the NIR and ALMA sources are unrelated, and the galaxy responsible for the submm emission may not be at  $z = 2.606$ .

and PYRAF.<sup>14</sup> Briefly, the processing consists of removing cosmic ray-like features, dividing by flat-fields taken with infrared and quartz halogen lamps, subtracting sky emission using exposures taken at a different point in the dither pattern, and finally rectifying the tilted, curved spectra using pinhole flats. Wavelength calibration is performed using Argon arc spectra, and then a spectrum of each order is extracted, divided by the standard star observation to cancel telluric absorption lines, and roughly flux-calibrated using the telluric standard star spectrum. The different spectral orders for each extraction window are merged into a single 1D spectrum from 0.84 to 2.48  $\mu\text{m}$ . In all cases, the agreement in flux between the overlapping regions of two consecutive orders was very good, and scaling factors of only 3 per cent or less were necessary.

### 2.1.3 XSHOOTER observations

One target was observed on 2015 March 4 with the XSHOOTER echelon spectrograph (Vernet et al. 2011) on the ESO VLT-UT2 (Kueyen) as part of programme 094.A-0811(A), providing near-continuous spectroscopy from 0.3 to 2.48  $\mu\text{m}$  with a 1.2 arcsec wide and 11 arcsec long slit. The slit was located on a possible NIR counterpart (but see footnote 15) located 2.5 arcsec from the ALMA source UDS 306.2, dithering the observations in an ABBA sequence at positions +3 and –3 arcsec along the slit axis. The observation was setup to first peak up on a nearby star in a field within 1 arcmin of the target position, and then a blind offset was performed. Twelve exposures of 300 s each were taken.

The ESO XSHOOTER pipeline (Modigliani et al. 2010) was used to reduce the data. This pipeline was used to perform spatial and spectral rectification on the spectra (which exhibited significant spatial curvature in addition to a non-linear wavelength scale) by using two-dimensional arc spectra obtained through a pinhole mask.

<sup>14</sup>PYRAF is a product of the Space Telescope Science Institute, which is operated by AURA for NASA.

For the IR channel, the data were mapped to an output spectral scale of  $1 \text{ \AA pix}^{-1}$  and a spatial scale of 0.21 arcsec (from original scales of approximately  $0.5 \text{ \AA pix}^{-1}$  and 0.24 arcsec, respectively). In the optical channel, the data were mapped to an output spectral scale of  $0.4 \text{ \AA pix}^{-1}$  and a spatial scale of 0.16 arcsec. In both channels, the data were flat-fielded, and cosmic rays were identified and masked. The two dither positions were subtracted to remove the sky to first order, and the different Echelle orders were combined together into a continuous spectrum (taking into account the variation in throughput with wavelength in different overlapping Echelle orders) before spatially registering and combining the data taken at the two dither positions and removing any residual sky background.

## 2.2 COSBO3

We also include COSBO3 (aka AzTEC-C6 and COSMOS 850.05), for which we obtained an  $H\alpha$ -based redshift for one component SMG in this work;  $H\alpha$ - and CO-based redshifts for additional component SMGs have been presented in the literature previously. COSBO3 was originally detected as a single millimetre source by MAMBO (Bertoldi et al. 2007) and AzTEC (Aretxaga et al. 2011) and at  $850 \mu\text{m}$  with SCUBA-2 (Casey et al. 2013). The accompanying  $450 \mu\text{m}$  maps from SCUBA-2 hinted at multiplicity, with the source being split into two independent sources at 7 arcsec resolution (identified in Casey et al. 2013 as COSMOS 450.16 and 450.28). Using ALMA, Bussmann et al. (2015) detected five-component SMGs. We observed component c with MOSFIRE using the same setup and reduced the data in the same manner as Casey et al. (2017). To complement the  $H\alpha$ -based spectroscopic redshifts, we also consider the CO-based spectroscopic redshifts for three of the component SMGs obtained in other work.

## 2.3 Single-dish submm sources drawn from the literature

We also analyse four single-dish submm sources from the literature, which, to the best of our knowledge, are the only other suitable multicomponent single-dish submm sources for which rest-frame UV/optical spectra are available for at least two of the component SMGs. The first two single-dish sources are GOODS 850-13 and 850-15; spectroscopic redshifts for some of the individual component SMGs that comprise these multicomponent single-dish sources were obtained in Swinbank et al. (2004), Wang et al. (2011), and Barger et al. (2012). The remaining two, ALESS 084 and 088, are from Danielson et al. (2017), who obtained spectra of many of the resolved ALESS SMGs (Hodge et al. 2013). Although taken at face value, the work of Danielson et al. (2017) contains seven multicomponent single-dish submm sources that meet our criteria (ALESS 017, 041, 067, 075, 080, 084, and 088), we include only two of them, ALESS 084 and 088, for the following reasons: we exclude ALESS 017, 075, and 080 because for each of these single-dish sources, one of the two component SMGs with spectroscopic redshifts is in the ‘supplementary’ sample of Hodge et al. (2013) owing to it lying outside the primary ALMA beam, and subsequent observations suggest that most of the ‘supplementary’ sources are artefacts of poor coverage of the  $u$ - $v$  plane (I. Smail, private communication). We exclude ALESS 041 because one of its two component SMGs has a spectroscopic redshift and the other’s spectrum exhibits continuum without features, but no photometric redshift is available for the second component. We exclude ALESS 067 because Danielson et al. (2017) claim that this single-dish submm source’s two component SMGs are at the

same redshift based on morphology, but they did not obtain two independent spectroscopic redshifts.

Furthermore, some of the component SMGs of the ALESS single-dish submm sources we do include, ALESS 084 and 088, have spectra of marginal quality:  $Q = 3$ , indicating redshifts that should not be considered fully independent of the photometric redshifts (and no spectroscopic redshifts based on  $Q = 3$  spectra from Danielson et al. 2017 have been subsequently confirmed; M. Swinbank, private communication), or  $Q = 4$ , corresponding to spectra that exhibit continuum but no sufficiently trustworthy features to assign a spectroscopic redshift (see Danielson et al. 2017 for further details). We do not include these components in the ‘unambiguous sample’, which is defined below.

## 2.4 Sub-sample definitions

Below, we analyse the ‘unambiguous’ sub-sample of single-dish submm sources, for which robust spectroscopic redshifts are available for at least two SMGs comprising a given single-dish submm source, separately. For the full sample, we additionally include component SMGs with less-robust spectroscopic redshifts (specifically, those assigned  $Q = 3$  by Danielson et al. 2017) and component SMGs whose spectra exhibit continuum but no features; for the latter, we employ photometric redshifts to compute the redshift separations of the components. The single-dish submm sources and component SMGs included in the two sub-samples (with the individual components specified in parentheses) are as follows:

(i) Unambiguous sub-sample: LOCK-03 (components a and b), LOCK-08 (a and b), UDS292 (0 and 1), COSBO3 (a–e), GOODS 850-13 (b and c), GOODS 850-15 (a and b), and ALESS 088 (1 and 5).

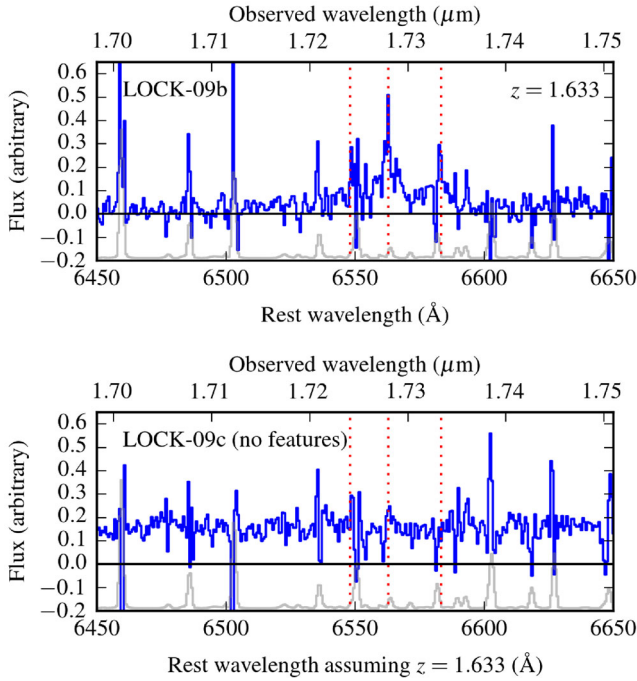
(ii) Full sample: All sources/component SMGs in the unambiguous sub-sample, plus LOCK-09 (b and c), UDS306 (0 and 1),<sup>15</sup> and ALESS 084 (1 and 2). We also include two additional component SMGs comprising ALESS 088 components 2 and 11 (components 1 and 5 were already included in the unambiguous sub-sample. Although Danielson et al. (2017) report spectroscopic redshifts for ALESS 088.2 and 11, they are based on marginal spectra ( $Q = 3$ ), so we opt to exclude these SMGs from the unambiguous sub-sample.

## 3 RESULTS

### 3.1 Discussion of individual sources

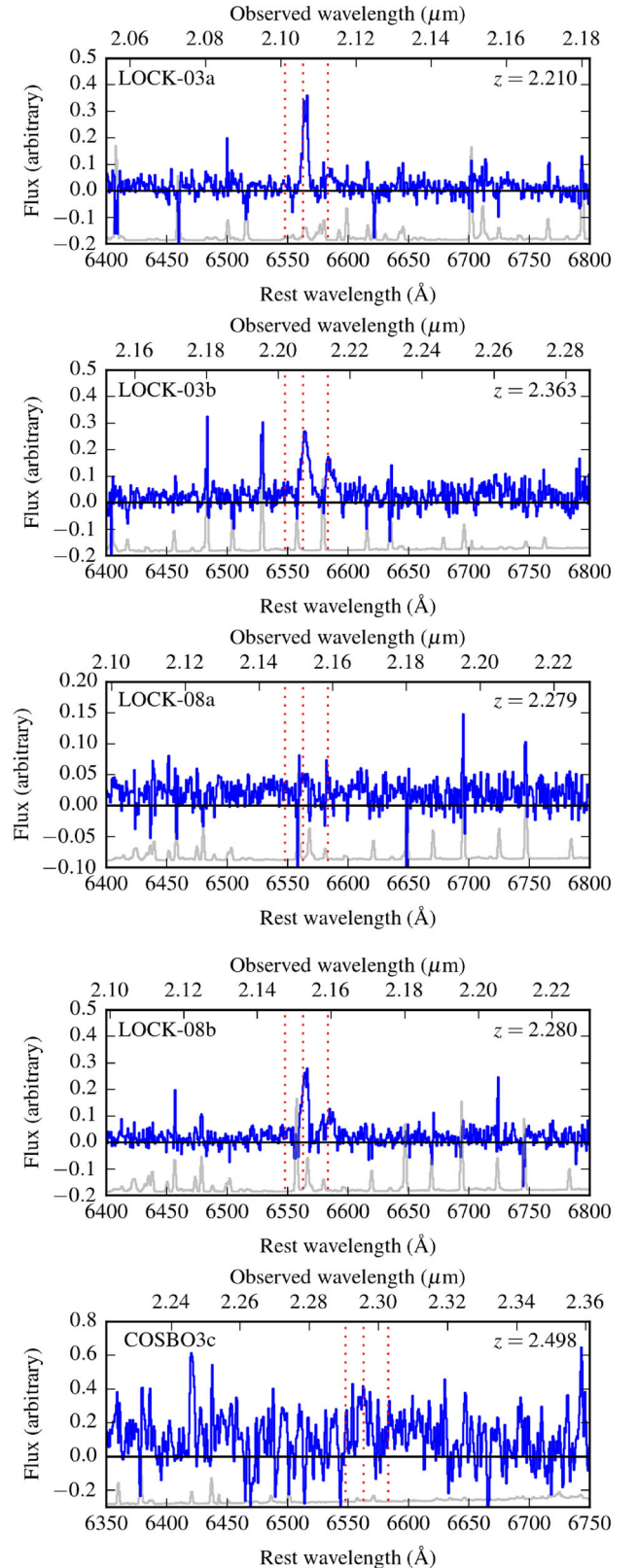
Fig. 1 shows example MOSFIRE  $H$ -band spectra of the two component SMGs of one of the single-dish submm sources in our sample, LOCK-09; the remaining spectra for the components of multicomponent single-dish submm sources obtained in this work using MOSFIRE, GNIRS, and XSHOOTER are shown in Figs 2, 3, and 4, respectively. Thumbnail images showing the interferometric submm sources are available in the works that presented the submm interferometric observations (Bussmann et al. 2015; Simpson et al. 2015a,b; Hill et al. 2017). Before discussing the contributions of associated and unassociated components, we first comment on the three single-dish submm sources for which the physical nature of the blending is somewhat ambiguous.

<sup>15</sup> As noted in Table 1, the large offset between the potential NIR counterpart to UDS306.2 and the ALMA source precludes us from assigning a redshift to UDS306.2. We present the XSHOOTER spectrum of the NIR counterpart for completeness, but we do not include UDS306.2 in the analysis.

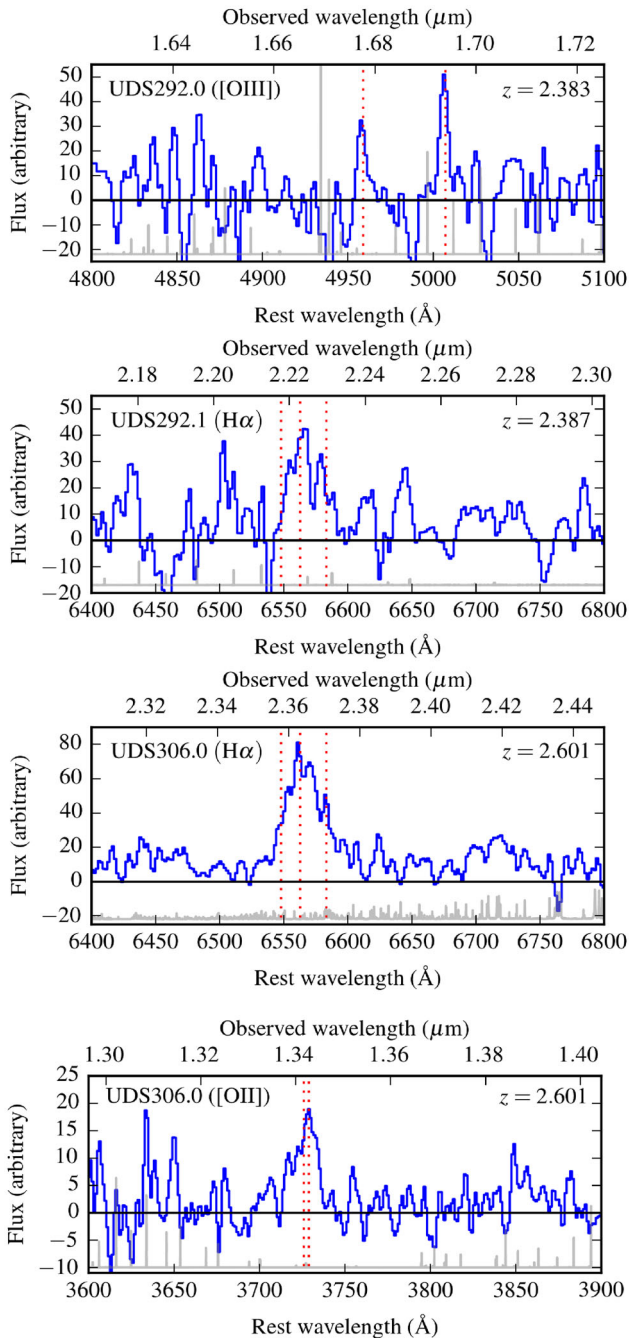


**Figure 1.** Sub-regions of the MOSFIRE  $H$ -band spectra of the NIR counterparts of LOCK-09b (top) and LOCK-09c (bottom). The grey lines show sky spectra (arbitrarily normalized and offset) from [https://www2.keck.hawaii.edu/inst/mosfire/sky\\_lines.html](https://www2.keck.hawaii.edu/inst/mosfire/sky_lines.html). The lines visible in the spectrum of LOCK-09b (marked with vertical dotted red lines) are identified as  $H\alpha$  and the  $[N\text{II}]$  doublet, yielding a firm spectroscopic redshift of 1.633. Strong continuum emission is detected from LOCK-09c, but  $H\alpha$  emission (at  $z = 1.633$ ) is absent. Although it is possible that the two component SMGs are at the same redshift, the most parsimonious interpretation is that the two SMGs are physically unassociated, as also suggested by LOCK-09c’s photometric redshift of  $0.90 \pm 0.05$ ; see the text for details.

$H\alpha$  and the  $[N\text{II}]$  doublet are detected in the spectrum of LOCK-09b (top panel of Fig. 1), yielding a firm spectroscopic redshift of 1.633. The spectrum of LOCK-09c (bottom panel) exhibits strong continuum emission but no obvious features. We argue that LOCK-09b and LOCK-09c are unlikely to be at the same redshift for the following reasons: LOCK-09b has an  $H\alpha$  equivalent width (EW) of  $120 \text{ \AA}$ ; if LOCK-09c is at the same redshift, its  $H\alpha$  EW is  $\lesssim 1 \text{ \AA}$ . For LOCK-09b,  $L_{H\alpha} = (5.65 \pm 0.20) \times 10^{-20} \text{ W m}^{-2}$  (a  $27\sigma$  detection), corresponding to an unobscured SFR of  $(5.6 \pm 0.2) M_{\odot} \text{ yr}^{-1}$ . For LOCK-09c, assuming the same  $z = 1.633$ , then  $L_{H\alpha} < 5 \times 10^{-21} \text{ W m}^{-2}$  ( $3\sigma$ ), corresponding to a  $3\sigma$  upper limit on the unobscured SFR of  $0.5 M_{\odot} \text{ yr}^{-1}$ . It is possible that these component SMGs are at the same redshift and that the nebular emission from LOCK-09c is too dust-obscured to be detected; however, this would require that LOCK-09c has an extremely low  $H\alpha$  equivalent width. For example, using a sample of 73 local (U)LIRGs, Poggianti & Wu (2000) obtained a minimum  $H\alpha + [N\text{II}]$  EW of  $17.7 \text{ \AA}$ . Moreover, given its submm flux density of  $S_{850} \approx 5 \text{ mJy}$ , LOCK-09c likely has an SFR in excess of  $500 M_{\odot} \text{ yr}^{-1}$  (e.g. da Cunha et al. 2015; Cowie et al. 2017); thus, in the scenario in which this component SMG is at  $z = 1.633$  and the  $H\alpha$  emission is simply too (differentially) dust-obscured to be detected,  $\text{SFR}_{H\alpha}/\text{SFR}_{\text{IR}} \lesssim 10^{-3}$  (i.e.  $A_{H\alpha} > 7.5$ ). Such a ratio would be low even for SMGs (e.g. Swinbank et al. 2004; Flores et al. 2004), but spectroscopic follow-up studies are likely biased against SMGs with the highest  $A_{H\alpha}$  values, and this scenario ( $A_{H\alpha} > 7.5$ ) is not impossible on these grounds alone. However,

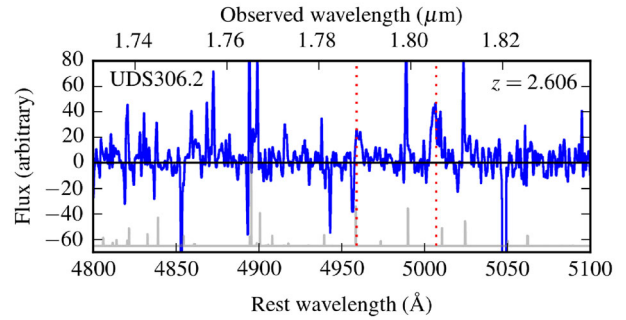


**Figure 2.** Sub-regions of the MOSFIRE spectra of the NIR counterparts of LOCK-03a, LOCK-03b, LOCK-08a, LOCK-08b, and COSBO3c near the features of interest. The grey lines show sky spectra (arbitrarily normalized and offset) from [https://www2.keck.hawaii.edu/inst/mosfire/sky\\_lines.html](https://www2.keck.hawaii.edu/inst/mosfire/sky_lines.html). The vertical dotted red lines denote the positions of the  $H\alpha$  and the  $[N\text{II}]$  doublet.



**Figure 3.** Sub-regions of the GNIRS spectra of the NIR counterparts of UDS292.0, UDS292.1, and UDS306.0 near the features of interest (labelled in the individual panels and marked with vertical dotted red lines). The grey lines show sky spectra (arbitrarily normalized and offset) obtained from the Gemini Observatory (Lord 1992; <http://www.gemini.edu/sciops/telescopes-and-sites/observing-condition-constraints/ir-background-spectra>).

given the detection of strong continuum emission, we consider the chance projection scenario to be more likely. Finally, LOCK-09c has  $z_{\text{phot}} = 0.90 \pm 0.05$ ; this differs from  $z = 1.633$  by  $14\sigma$ , which further supports our interpretation that LOCK-09c is not at  $z = 1.633$  (the photometric redshift of LOCK-09b,  $z_{\text{phot}} = 1.46 \pm 0.19$ , is consistent with its spectroscopic redshift). We thus conclude that the most likely explanation for the detection of continuum only



**Figure 4.** Sub-region of the XSHOOTER spectrum of the potential NIR counterpart of UDS306.2 near the [O III] emission lines at 4959 and 5007 Å (marked by the vertical dotted red lines). The grey line shows a sky spectrum (arbitrarily normalized and offset) obtained from the Gemini Observatory (Lord 1992; <http://www.gemini.edu/sciops/telescopes-and-sites/observing-condition-constraints/ir-background-spectra>).

from LOCK-09c is that it is a chance projection, but obtaining a spectroscopic redshift for LOCK-09b is necessary to definitively confirm this conclusion.

For UDS306.1, we detect continuum but no obvious features. However, we do not interpret this component as a chance projection because its photometric redshift is discrepant from the spectroscopic redshift of UDS306.0 by  $<2.5\sigma$ . Moreover, the upper limit on the unobscured SFR ( $\lesssim 63 M_{\odot} \text{ yr}^{-1}$  at  $3\sigma$ ) is weaker than that for LOCK-09c, so we cannot appeal to the arguments we used above.

The sample drawn from the literature also contains one single-dish submm source for which continuum only is detected from one component SMG (ALESS 084). Because H $\alpha$  luminosities are not reported in Danielson et al. (2017), we cannot make SFR-based arguments similar to those made for LOCK-09 above. For this single-dish source, the photometric redshift of the continuum-only component differs from the spectroscopic redshift of the other component by  $>27\sigma$ . However, for the component with a spectroscopic redshift, the photometric redshift is discrepant from the spectroscopic redshift by almost  $23\sigma$ , likely because the photometric redshift uncertainties quoted by Simpson et al. (2014) are unrealistically small, but also perhaps because the spectroscopic redshifts based on marginal-quality ( $Q = 3$ ) spectra from Danielson et al. (2017) are not trustworthy. We thus label this single-dish source as a ‘likely projection’, but future observations may reveal that the two component SMGs are actually associated. Regardless, we note that excluding ALESS 084 does not change our qualitative conclusions.

### 3.2 Statistical constraints

For each of the multicomponent single-dish submm sources in the unambiguous sample, we compute the difference in redshift between the components for which we have spectroscopic redshifts,  $\Delta z$ ; for single-dish sources with more than two component SMGs, following Hayward et al. (2013b), we compute the redshift separation of each subdominant component from the brightest component and add these separations in quadrature. For the full sample, we compute  $\Delta z$  using spectroscopic redshifts when available and photometric redshifts only for component SMGs without spectroscopic redshifts.<sup>16</sup> We then classify the physical nature of the

<sup>16</sup> The  $\Delta z$  values based on photometric redshifts should be interpreted with caution because obtaining accurate photometric redshifts for dust-obscured galaxies is challenging. For many SMGs, the spectroscopic redshifts are formally significantly discrepant from the photometric redshifts reported in

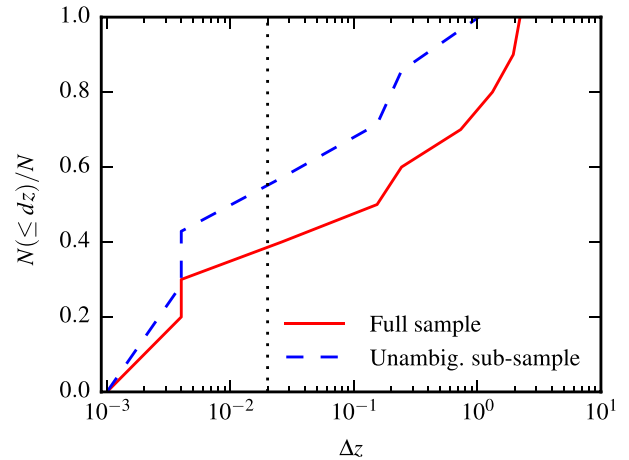
multiplicity of each single-dish submm source using the criterion defined in Hayward et al. (2013b), which was selected because this value separates the two peaks of the bimodal  $\Delta z$  distribution predicted by their theoretical model (see their fig. 4): associated sources are those with  $\Delta z \leq 0.02$ , whereas those with greater  $\Delta z$  values are considered chance projections of at least one unassociated component SMG. Naïvely, considering only the velocity difference between two galaxies and asking whether the galaxies are bound, a threshold of  $\Delta v \gtrsim 1000 \text{ km s}^{-1}$ , or  $\Delta z \gtrsim 0.003$ , would be sufficient to identify unbound – and thus non-merging – pairs, and an even lower threshold could likely be used. However, we also wish to classify unbound but still associated pairs, such as those contained in the same ‘proto-cluster’ or dark matter filament, as physically associated. Additionally, for single-dish sources with more than two component SMGs, the fact that we add the pairwise redshift separations in quadrature means that the velocity difference inferred from the  $\Delta z$  value will be greater than the velocity differences between individual galaxies. Moreover, because the distribution of  $\Delta z$  in predicted by theoretical models (Hayward et al. 2013b; Cowley et al. 2015; Muñoz Arancibia et al. 2015) is strongly bimodal, using a somewhat smaller or larger threshold does not affect the model predictions. However, one of our single-dish sources, COSBO3, has a  $\Delta z$  value very close to this threshold,  $\Delta z = 0.026$ . Although we formally treat this as a chance projection to ensure the fairest possible comparison with the model predictions, we note that this classification is sensitive to the exact threshold employed, and this association of SMGs has been previously interpreted as a ‘proto-cluster’ (Casey et al. 2015; Wang et al. 2016).

In Fig. 5, we show cumulative distribution functions (CDFs) of  $\Delta z$  for the two sub-samples defined above; the redshift separations and nature of the component multiplicity are presented in Table 2. The dashed blue line represents the CDF of the unambiguous sample (i.e. only robust spectroscopic redshifts are used). The solid red line shows the CDF of the full sample. The CDFs reveal that for both sub-samples, of order half the single-dish submm sources are chance projections of at least one unassociated component (i.e. have  $\Delta z > 0.02$ ). For the unambiguous sub-sample and the full sample, the mean  $\Delta z$  values are  $0.21 \pm 0.05$  and  $0.67 \pm 0.23$ , respectively.

Of the seven single-dish submm sources in the unambiguous sub-sample, in three sources (3/7, or  $43^{+39}_{-33}$  per cent)<sup>17</sup> – LOCK-08, UDS292, and GOODS 850-15 – the two component SMGs for which spectroscopic redshifts are available are physically associated. Interestingly, because the components of LOCK-08 have a projected separation of  $\sim 25$  arcsec ( $\sim 200$  kpc), they are unlikely to be undergoing a merger but rather simply part of the same dark matter filament. The remaining four single-dish sources (4/7, or  $57^{+33}_{-39}$  per cent) in the unambiguous sub-sample (LOCK-03, COSBO3, GOODS 850-13, and ALESS 088) are classified as chance projections of at least one unassociated SMG, although as already noted above, COSBO3 has a  $\Delta z$  value only slightly greater than the threshold for a chance projection; given the significant

other works. These discrepancies suggest that the original works underestimated the uncertainties on their photometric redshifts. However, even if the uncertainties on the photometric redshifts were multiplied by a factor of a few to account for this likely underestimation, the  $\Delta z$  values computed using photometric redshifts for LOCK-09, ALESS 084, and ALESS 088 would remain significantly greater than 0.02, and thus their classification as chance projections should be robust to this issue.

<sup>17</sup> Throughout the work, the quoted uncertainties on percentages correspond to 95 per cent binomial confidence intervals (Clopper & Pearson 1934).



**Figure 5.** Empirical CDFs of  $\Delta z$  for the multicomponent single-dish submm sources in our sample (including those drawn from the literature). The blue dashed line represents the CDF for the unambiguous sample (i.e. when only robust spectroscopic redshifts are used), whereas the solid red line shows the CDF for the full sample (for components whose spectra exhibit continuum only, photometric redshifts are used to compute  $\Delta z$ ). The vertical dotted line represents the separation between associated and unassociated single-dish submm sources  $\Delta z = 0.02$ ; this value is motivated by the bimodal  $\Delta z$  distributions predicted by multiple theoretical models (Hayward et al. 2013b; Cowley et al. 2015; Muñoz Arancibia et al. 2015). In both sub-samples, of order half of the single-dish submm sources contain at least one unassociated SMG.

**Table 2.** Redshift and angular separations of the components and nature of the multiplicity of single-dish submm sources.

Single-dish ID	$\Delta z^a$	Angular separation <sup>b</sup> (arcsec)	Nature of multiplicity
LOCK-03	0.154	11.3	Unassociated
LOCK-08	0.001	24.7	Associated
LOCK-09	$(0.73 \pm 0.05)$	$(12.0)$	Unassociated
UDS292	0.004	6.7	Associated
UDS306	$(1.32^{+0.06}_{-0.53})$	$(6.1)$	Ambiguous
COSBO3 <sup>c</sup>	0.026	18.1	Unassociated
GOODS 850-13	0.243	16.4	Unassociated
GOODS 850-15	0.004	2.0	Associated
ALESS 084	$(2.22^{+0.19}_{-0.08})$	8.7	Unassociated
ALESS 088	$1.026 (1.951)$	$15.0 (17.9)$	Unassociated

<sup>a</sup>Redshift separation of components (see the text for definition). Values in parentheses are based on one or more photometric redshifts (used only for component SMGs whose spectra exhibit continuum but no features; for these, error bars are quoted for the  $\Delta z$  values) or less-robust ( $Q = 3$ ) spectroscopic redshifts from Danielson et al. (2017).

<sup>b</sup>Angular separation of the components used to compute  $\Delta z$ ; when there are more than two component SMGs, the pairwise angular separations are added in quadrature in the same manner as for  $\Delta z$ . For reference, in the redshift range of interest,  $z \sim 1-4$ , 1 arcsec corresponds to 7–8 kpc.

<sup>c</sup>Although we classify COSBO3 as a chance projection because it has  $\Delta z = 0.026 > 0.02$ , we note that this  $\Delta z$  value is very close to the adopted threshold, and this source has been interpreted as a ‘proto-cluster’ by other authors (Casey et al. 2015; Wang et al. 2016).

uncertainties in the associated and unassociated fractions as a result of the small sample size, classifying COSBO3 as physically associated or removing it entirely would not materially affect our conclusions. LOCK-03 is particularly interesting because the MOS-FIRE spectrum of one of the resolved SMGs, LOCK-03a, exhibits two kinematically distinct components separated by 0.8 arcsec



( $\sim 7$  kpc) and  $160 \text{ km s}^{-1}$ . Thus, it appears that the ALMA source LOCK-03a corresponds to a late-stage, pre-coalescence merger (i.e. is a ‘galaxy-pair SMG’ in the parlance of Hayward et al. 2011), and it may be resolved into two separate sources in future higher resolution interferometric submm observations. The other ALMA component, LOCK-03b, which has a submm flux density equal to that of LOCK-03a, is at a significantly different redshift (the two sources are separated by  $\Delta z = 0.154$ ). LOCK-03 thus appears to be a blend of an ongoing merger and an unassociated SMG.

Turning to the full sample (i.e. using photometric redshifts to compute  $\Delta z$  for component SMGs from which continuum only was detected), an additional three single-dish submm sources (LOCK-09, UDS306, and ALESS 084) would naïvely be classified as chance projections based on their having  $\Delta z > 0.02$ . However, regarding UDS306, because the photometric redshifts of the two component SMGs are discrepant by  $< 2.5\sigma$ , we consider the nature of the blending unconstrained. Thus, of the nine single-dish sources in the full sample for which we have firm or tentative constraints on the nature of the multiplicity, six (6/9, or  $67_{-37}^{+26}$  per cent) are comprised of at least one unassociated component SMG. The full sample thus contains a greater fraction of chance projections than the unambiguous sub-sample, but given the large uncertainties, the difference is not statistically significant.

#### 4 SUMMARY AND DISCUSSION

We have obtained spectroscopic constraints on the redshifts of individual components of six multicomponent single-dish submm sources to investigate the nature of single-dish submm source multiplicity (i.e. whether the component SMGs are physically associated or chance projections). We supplemented our sample with four single-dish sources from the literature. Of the seven multicomponent single-dish submm sources in our sample for which robust spectroscopic redshifts are available for at least two component SMGs, only three (3/7, or  $43_{-33}^{+39}$  per cent) are clearly blends of physically associated (but not necessarily merging) SMGs. Considering also constraints based on detection of continuum but no features, for which photometric redshifts are used to calculate the redshift separation, and spectroscopic redshifts based on marginal spectra, the nature of the multiplicity of nine of the single-dish sources can be tentatively constrained. Of these, six (6/9, or  $67_{-37}^{+26}$  per cent) are comprised of at least one unassociated component SMG. The nature of the multiplicity of one single-dish source, UDS 306, is unconstrained by the available data.

This work is the first to place statistical constraints on the relative contributions of physically associated SMGs and chance projections to the multicomponent single-dish submm source population. Such constraints can help distinguish amongst competing theoretical models for the SMG population. For example, in models in which mergers dominate the submm counts, the components of multicomponent SMGs should correspond to pre-coalescence mergers (Hayward et al. 2011, 2012, 2013a) and thus be close in terms of both redshift and angular separation. However, if late-stage merger-induced starbursts dominate the single-dish submm source population, high multiplicity would not be expected because such sources would not be resolved into multiple component SMGs. At the least, the fact that we observe a significant fraction of chance projections – in addition to single-dish submm sources comprised of associated SMGs that are widely separated – demonstrates that the classical view of bright submm sources as predominantly merger-induced starbursts (e.g. Engel et al. 2010) is incomplete. Put oth-

erwise, the combined effect of mergers – i.e. merger-induced starbursts and blending of pre-coalescence merging galaxies can both result in bright single-dish submm sources (Hayward et al. 2013a) – alone is insufficient to explain the observed submm counts if chance projections of unassociated SMGs are common amongst the single-dish submm source population, as our results suggest.

Although few theoretical models of single-dish submm sources have treated blending of physically unassociated SMGs, let alone chance projections, both Hayward et al. (2013b) and Cowley et al. (2015) concluded that the majority of multicomponent single-dish submm sources should be comprised of at least one physically unassociated SMG. However, the quantitative predictions of these models differ: whereas Cowley et al. (2015) predict that an almost-negligible fraction of multicomponent single-dish submm sources are comprised of solely physically associated SMGs (see their fig. 8), in the Hayward et al. (2013b) model, such sources account for  $\sim 15$  per cent of the population. The fraction of single-dish sources comprised of at least one unassociated component SMG found in this work ( $57_{-39}^{+33}$  and  $67_{-37}^{+26}$  per cent for the unambiguous sub-sample and the full sample, respectively) is more consistent with the predictions of Hayward et al. (2013b) than those of Cowley et al. (2015). However, the model predictions are sensitive to details such as the detection limits of the single-dish and interferometric observations considered. Moreover, the sample here may be subject to various biases and is modest in size. For these reasons, we caution against overinterpreting this comparison, and we defer a detailed comparison with models to future work. Nevertheless, our results qualitatively support the predictions of the aforementioned theoretical works, and future studies of larger samples should help distinguish amongst these and other theoretical models intended to reproduce the submm number counts.

We end with some caveats. In addition to being modest in size, the present sample is likely to be biased. Regarding our own observations, five of the six single-dish sources were selected from amongst the brightest SMGs in the S2CLS, and it is expected that the contribution of chance projections depends on the single-dish flux (Hayward et al. 2013b; Cowley et al. 2015). Moreover, our SMA data do not have sufficient resolution to distinguish mergers near coalescence (with projected separation  $\lesssim 2$  arcsec, or  $\lesssim 15$  kpc at the relevant redshifts), which implies a bias against physically associated multiples; the ALMA data are affected by the same bias but to a lesser extent. Regarding the sample drawn from the literature, no uniform selection was applied, and data censoring is likely to be an issue. Despite the above caveats, our work demonstrates that the contribution of chance projections to the single-dish submm source population cannot be ignored and provides motivation for constraining the physical nature of single-dish submm source multiplicity using a significantly larger, uniformly selected sample in future work.

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## REFERENCES

Aretxaga I. et al., 2011, *MNRAS*, 415, 3831  
 Astropy Collaboration, 2013, *A&A*, 558, A33  
 Barger A. J., Cowie L. L., Wang W.-H., 2008, *ApJ*, 689, 687  
 Barger A. J., Wang W.-H., Cowie L. L., Owen F. N., Chen C.-C., Williams J. P., 2012, *ApJ*, 761, 89  
 Barger A. J. et al., 2014, *ApJ*, 784, 9  
 Bertoldi F. et al., 2007, *ApJS*, 172, 132  
 Blain A. W., Smail I., Ivison R. J., Kneib J.-P., Frayer D. T., 2002, *Phys. Rep.*, 369, 111  
 Bluck A. F. L., Conzelmann C. J., Buitrago F., Grützbach R., Hoyos C., Mortlock A., Bauer A. E., 2012, *ApJ*, 747, 34  
 Bussmann R. S. et al., 2015, *ApJ*, 812, 43  
 Casey C. M. et al., 2013, *MNRAS*, 436, 1919  
 Casey C. M., Narayanan D., Cooray A., 2014, *Phys. Rep.*, 541, 45  
 Casey C. M. et al., 2015, *ApJ*, 808, L33  
 Casey C. M. et al., 2017, *ApJ*, 840, 101  
 Chapman S. C., Blain A. W., Smail I., Ivison R. J., 2005, *ApJ*, 622, 772  
 Chen C.-C., Cowie L. L., Barger A. J., Casey C. M., Lee N., Sanders D. B., Wang W.-H., Williams J. P., 2013, *ApJ*, 776, 131  
 Clopper C. J., Pearson E. S., 1934, *Biometrika*, 26, 404  
 Cowie L. L., Barger A. J., Hsu L.-Y., Chen C.-C., Owen F. N., Wang W.-H., 2017, *ApJ*, 837, 139  
 Cowley W. I., Lacey C. G., Baugh C. M., Cole S., 2015, *MNRAS*, 446, 1784  
 da Cunha E. et al., 2015, *ApJ*, 806, 110  
 Danielson A. L. R. et al., 2017, *ApJ*, 840, 78

Engel H. et al., 2010, *ApJ*, 724, 233  
 Flores H., Hammer F., Elbaz D., Cesarsky C. J., Liang Y. C., Fadda D., Gruel N., 2004, *A&A*, 415, 885  
 Geach J. E. et al., 2017, *MNRAS*, 465, 1789  
 Hayward C. C., Kereš D., Jonsson P., Narayanan D., Cox T. J., Hernquist L., 2011, *ApJ*, 743, 159  
 Hayward C. C., Jonsson P., Kereš D., Magnelli B., Hernquist L., Cox T. J., 2012, *MNRAS*, 424, 951  
 Hayward C. C., Narayanan D., Kereš D., Jonsson P., Hopkins P. F., Cox T. J., Hernquist L., 2013a, *MNRAS*, 428, 2529  
 Hayward C. C., Behroozi P. S., Somerville R. S., Primack J. R., Moreno J., Wechsler R. H., 2013b, *MNRAS*, 434, 2572  
 Hill R. et al., 2017, preprint (arXiv:1710.02231)  
 Hodge J. A. et al., 2013, *ApJ*, 768, 91  
 Hughes D. H. et al., 1998, *Nature*, 394, 241  
 Ivison R. J. et al., 2002, *MNRAS*, 337, 1  
 Ivison R. J. et al., 2007, *MNRAS*, 380, 199  
 Karim A. et al., 2013, *MNRAS*, 432, 2  
 Koprowski M. P., Dunlop J. S., Michałowski M. J., Cirasuolo M., Bowler R. A. A., 2014, *MNRAS*, 444, 117  
 Laigle C. et al., 2016, *ApJS*, 224, 24  
 Lord S. D., 1992, Technical report, A new software tool for computing Earth's atmospheric transmission of near- and far-infrared radiation, NASA Technical Memorandum 103957  
 McLean I. S. et al., 2010, in McLean I. S., Ramsay S. K., Takami H., eds, Proc. SPIE Conf. Ser. Vol. 7735, Ground-based and Airborne Instrumentation for Astronomy III. SPIE, Bellingham, p. 77351E  
 McLean I. S. et al., 2012, in McLean I. S., Ramsay S. K., Takami H., eds, Proc. SPIE Conf. Ser. Vol. 8446, Ground-based and Airborne Instrumentation for Astronomy IV. SPIE, Bellingham, p. 84460J  
 Michałowski M. J. et al., 2017, *MNRAS*, 469, 492  
 Modigliani A. et al., 2010, in Silva D. R., Peck A. B., Soifer B. T., eds, Proc. SPIE Conf. Ser. Vol. 7737, Observatory Operations: Strategies, Processes, and Systems III. SPIE, Bellingham, p. 773728  
 Muñoz Arancibia A. M., Navarrete F. P., Padilla N. D., Cora S. A., Gawiser E., Kurczynski P., Ruiz A. N., 2015, *MNRAS*, 446, 2291  
 Narayanan D. et al., 2015, *Nature*, 525, 496  
 Ochsenbein F., Bauer P., Marcout J., 2000, *A&AS*, 143, 23  
 Planck Collaboration XIII, 2016, *A&A*, 594, A13  
 Poggianti B. M., Wu H., 2000, *ApJ*, 529, 157  
 Roseboom I. G. et al., 2012, *MNRAS*, 419, 2758  
 Simpson J. M. et al., 2014, *ApJ*, 788, 125  
 Simpson J. M. et al., 2015a, *ApJ*, 799, 81  
 Simpson J. M. et al., 2015b, *ApJ*, 807, 128  
 Simpson J. M. et al., 2017, *ApJ*, 839, 58  
 Smith D. J. B., Hayward C. C., Jarvis M. J., Simpson C., 2017, *MNRAS*, 471, 2453  
 Smolčić V. et al., 2012, *A&A*, 548, A4  
 Steidel C. C. et al., 2014, *ApJ*, 795, 165  
 Strazzullo V., Pannella M., Owen F. N., Bender R., Morrison G. E., Wang W.-H., Shupe D. L., 2010, *ApJ*, 714, 1305  
 Strom A. L., Steidel C. C., Rudie G. C., Trainor R. F., Pettini M., Reddy N. A., 2017, *ApJ*, 836, 164  
 Swinbank A. M., Smail I., Chapman S. C., Blain A. W., Ivison R. J., Keel W. C., 2004, *ApJ*, 617, 64  
 Trouille L., Barger A. J., Cowie L. L., Yang Y., Mushotzky R. F., 2008, *ApJS*, 179, 1  
 Vernet J. et al., 2011, *A&A*, 536, A105  
 Wang W.-H., Cowie L. L., Barger A. J., 2006, *ApJ*, 647, 74  
 Wang W.-H., Cowie L. L., Barger A. J., Williams J. P., 2011, *ApJ*, 726, L18  
 Wang T. et al., 2016, *ApJ*, 828, 56  
 Wiklind T. et al., 2014, *ApJ*, 785, 111  
 Younger J. D. et al., 2009, *ApJ*, 704, 803  
 Zavala J. A. et al., 2015, *MNRAS*, 452, 1140

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